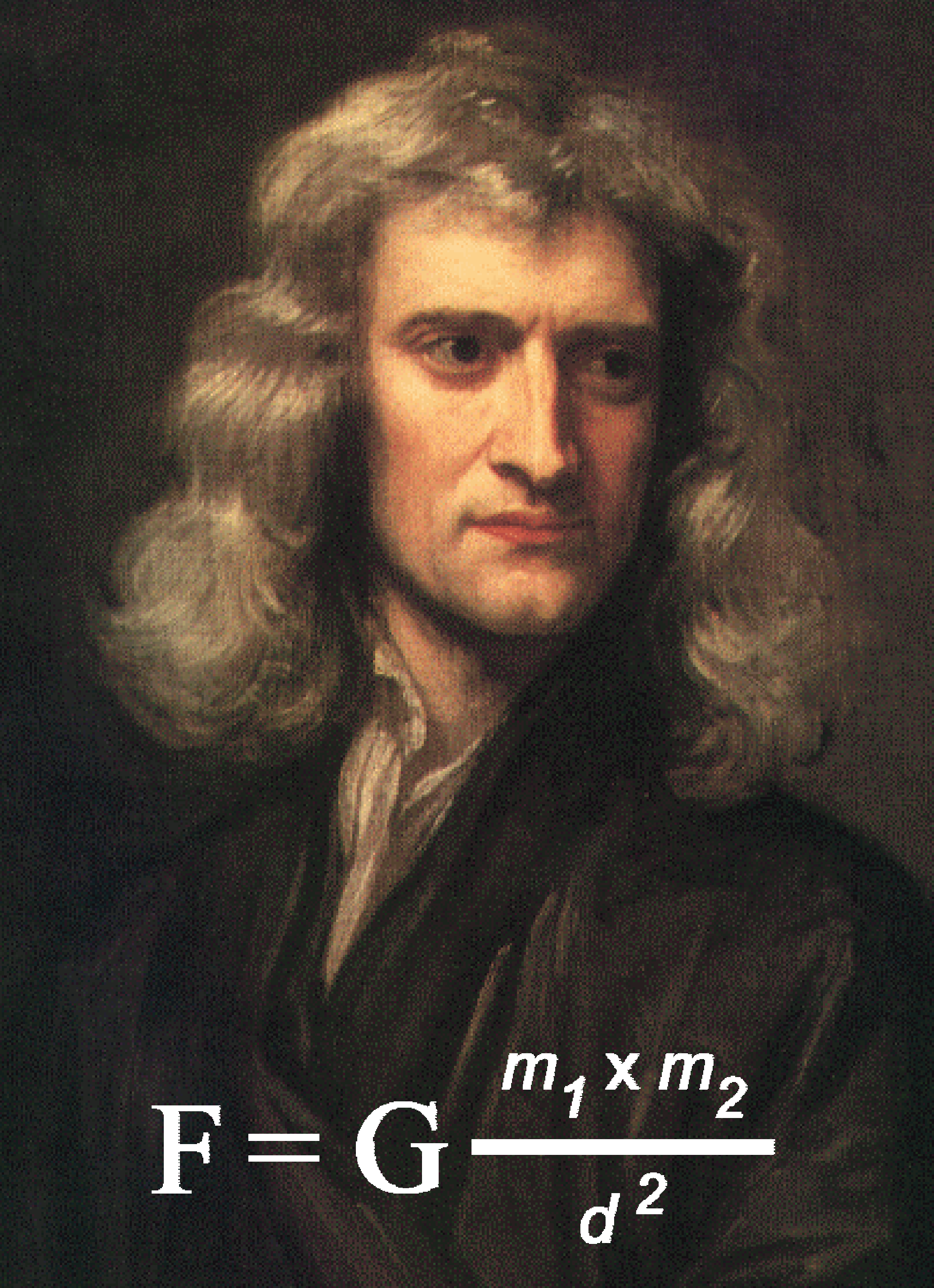
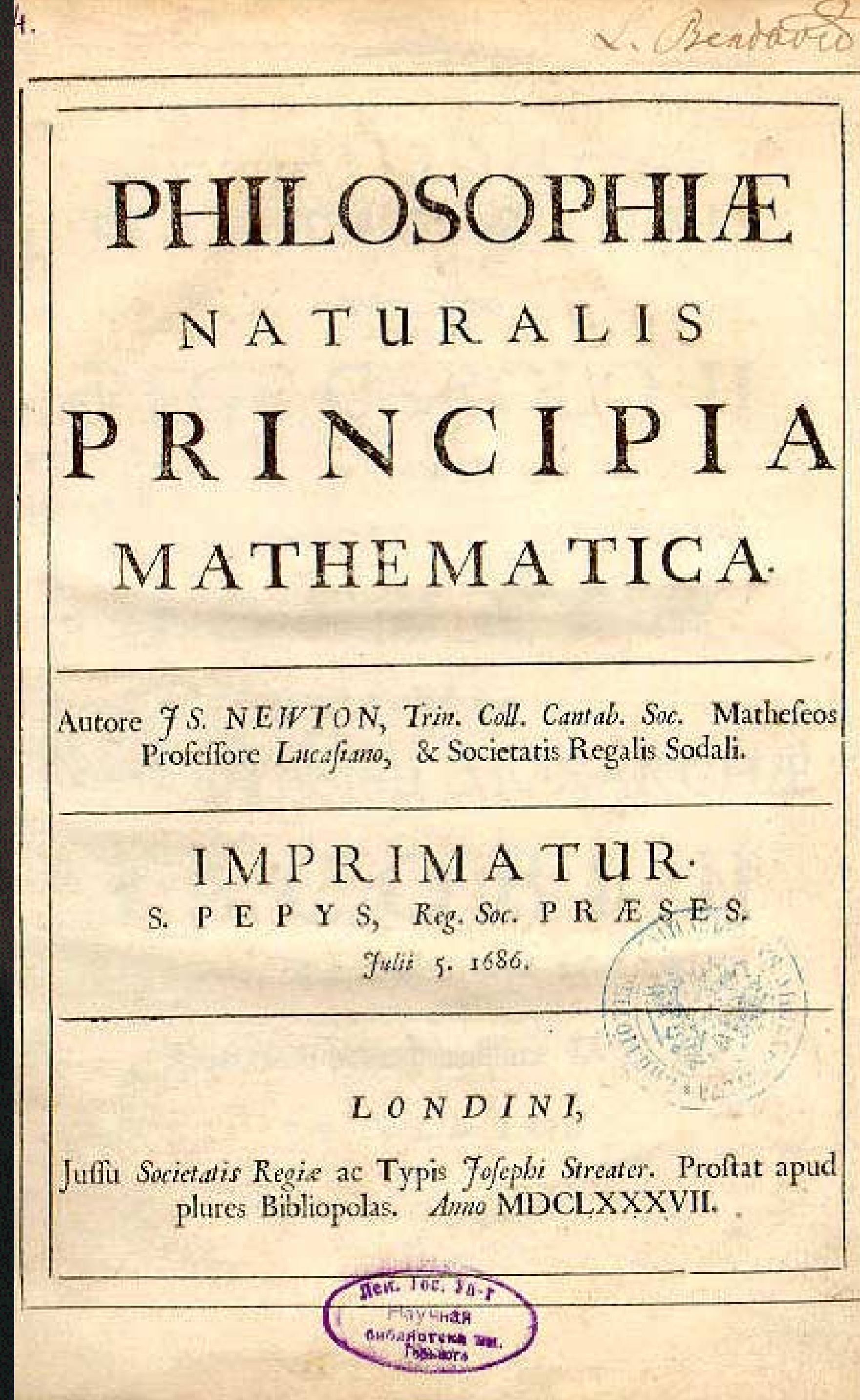


$$F = G\frac{m_1 \times m_2}{d^2}$$



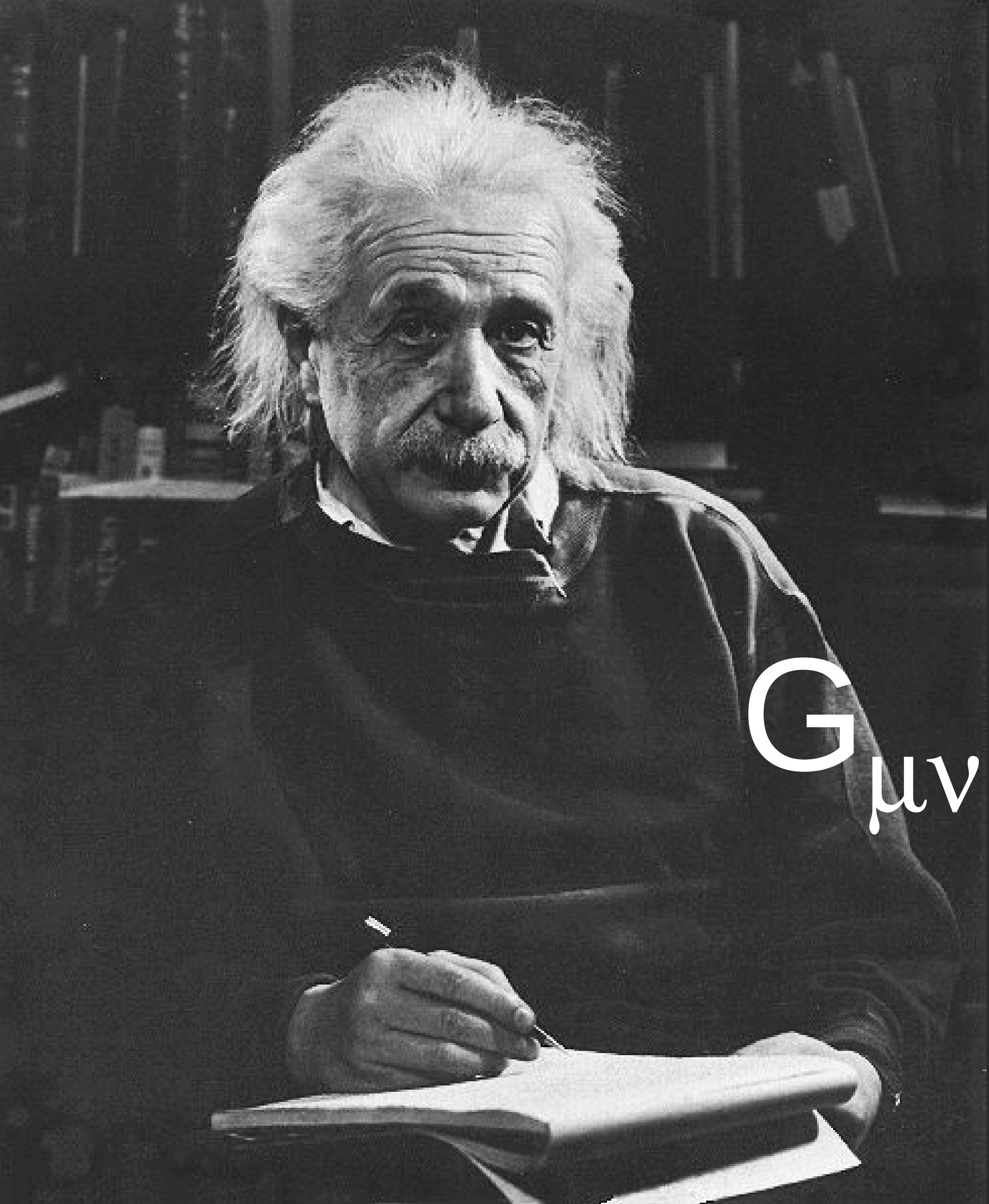
PHILOSOPHIÆ NATURALIS PRINCIPIA MATHEMATICA.

Autore *JS.* NEWTON, *Trin. Coll. Cantab. Soc.* Matheseos Professore *Lucasiano*, & Societatis Regalis Sodali.

IMPRIMATUR.
S. PEPYS, *Reg. Soc.* PRÆSES.
Julii 5. 1686.

LONDINI,

Jussu *Societatis Regiæ* ac Typis *Josephi Streater*. Prostat apud plures Bibliopolas. *Anno* MDCLXXXVII.

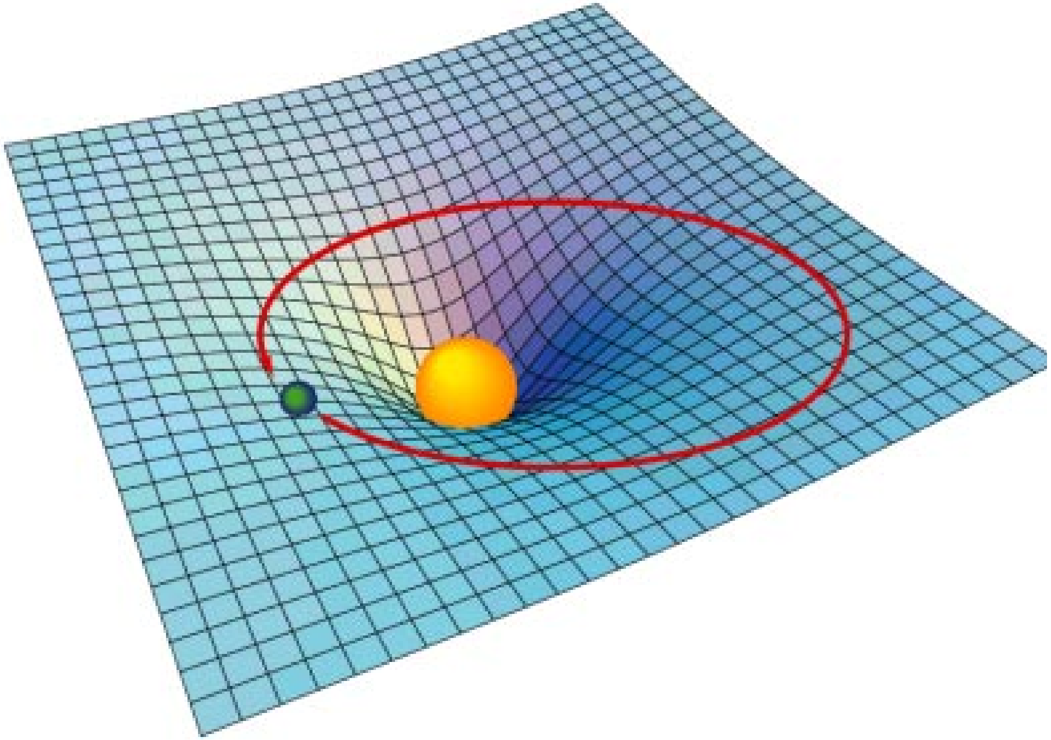


$$G_{\mu\nu} = 8\pi T_{\mu\nu}$$

LIGO

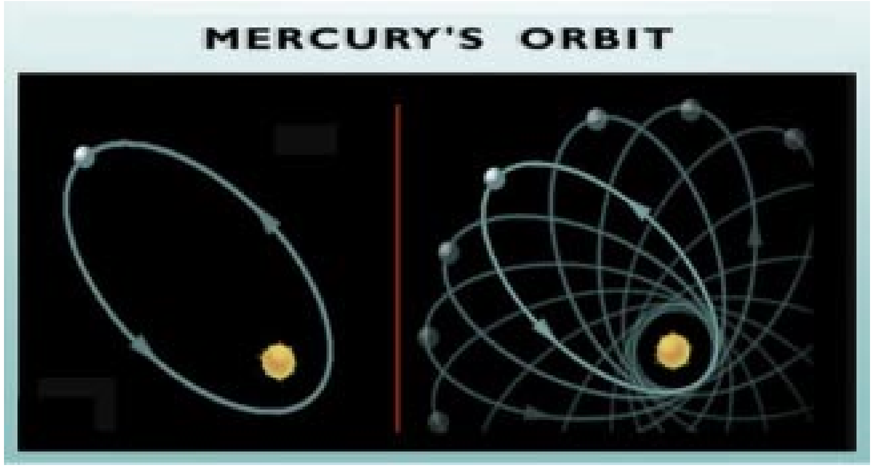
General Relativity

Einstein theorized that smaller masses travel toward larger masses, not because they are "attracted" by a mysterious force, but because the smaller objects travel through space that is warped by the larger object



- **Imagine space as a stretched rubber sheet.**
- **A mass on the surface will cause a deformation.**
- **Another mass dropped onto the sheet will roll toward that mass.**

LIGO Einstein's Theory of Gravitation
experimental tests



Mercury's orbit
perihelion shifts forward an extra +43"/century compared to Newton's theory

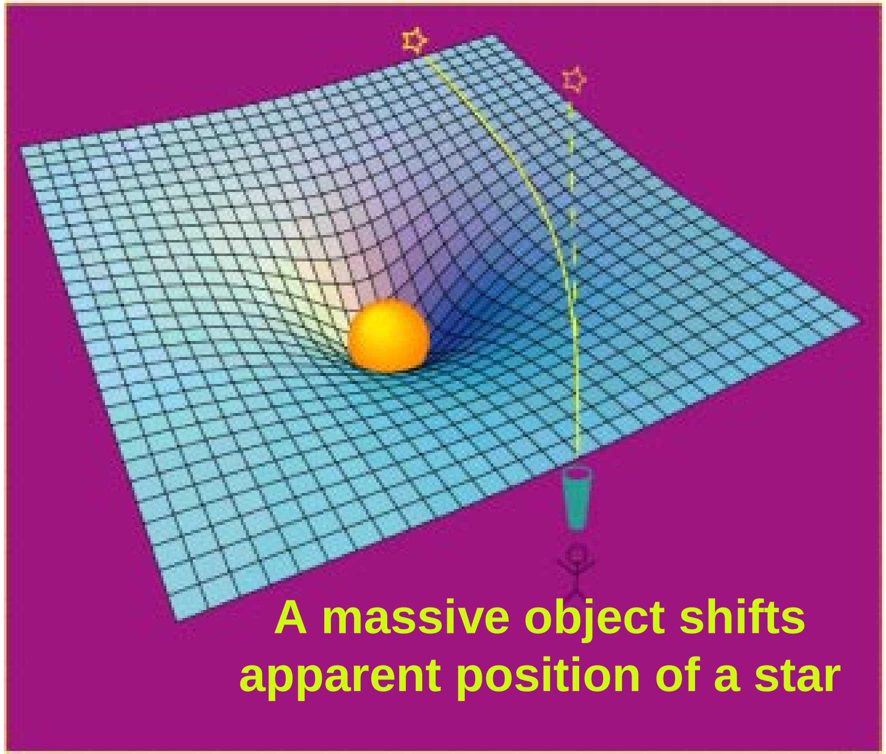
Mercury's elliptical path around the Sun shifts slightly with each orbit such that its closest point to the Sun (or "perihelion") shifts forward with each pass.

Astronomers had been aware for two centuries of a small flaw in the orbit, as predicted by Newton's laws.

Einstein's predictions exactly matched the observation.

New Wrinkle on Equivalence
bending of light

- **Not only the path of matter, but even the path of light is affected by gravity from massive objects**



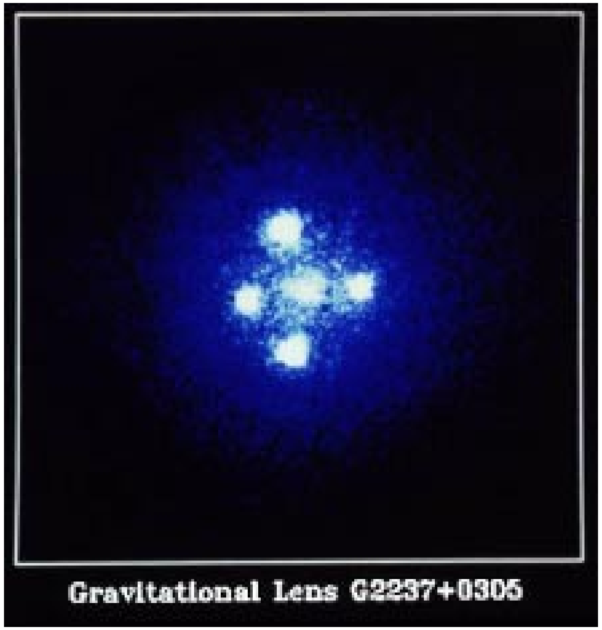
- **First observed during the solar eclipse of 1919 by Sir Arthur Eddington, when the Sun was silhouetted against the Hyades star cluster**
- **Their measurements showed that the light from these stars was bent as it grazed the Sun, by the exact amount of Einstein's predictions.**

The light never changes course, but merely follows the curvature of space. Astronomers now refer to this displacement of light as gravitational lensing.

LIGO Einstein's Theory of Gravitation

experimental tests

"Einstein Cross"
The bending of light rays
gravitational lensing



Quasar image appears around the central glow formed by nearby galaxy. The Einstein Cross is only visible in southern hemisphere.

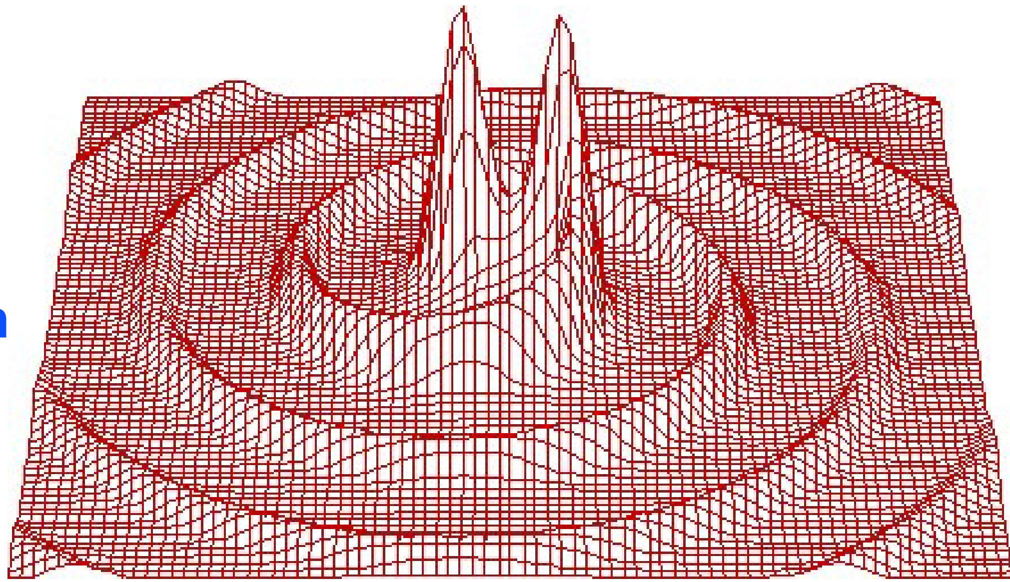
In modern astronomy, such gravitational lensing images are used to detect a 'dark matter' body as the central object

LIGO **Einstein's Theory of Gravitation**

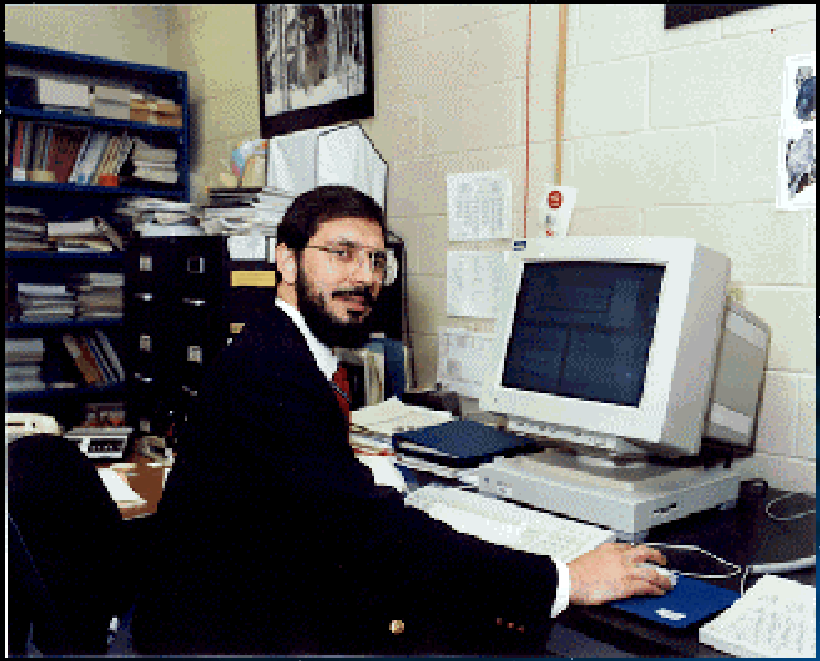
gravitational waves

- **a necessary consequence of Special Relativity with its finite speed for information transfer**

- **time dependent gravitational fields come from the acceleration of masses and propagate away from their sources as a space-time warpage at the speed of light**



gravitational radiation
binary inspiral of compact objects



Russel A. Hulse

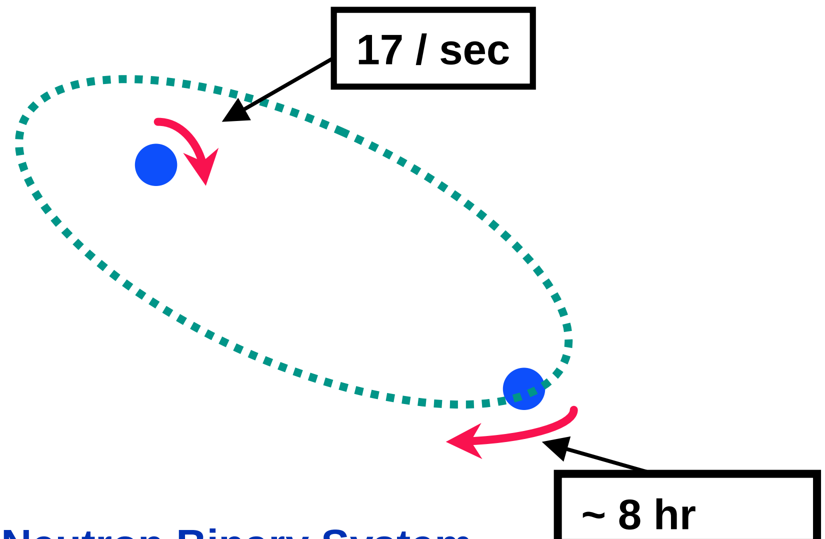
Joseph H.Taylor Jr

Source: www.NSF.gov

Gravitational Waves
the evidence

Neutron Binary System – Hulse & Taylor

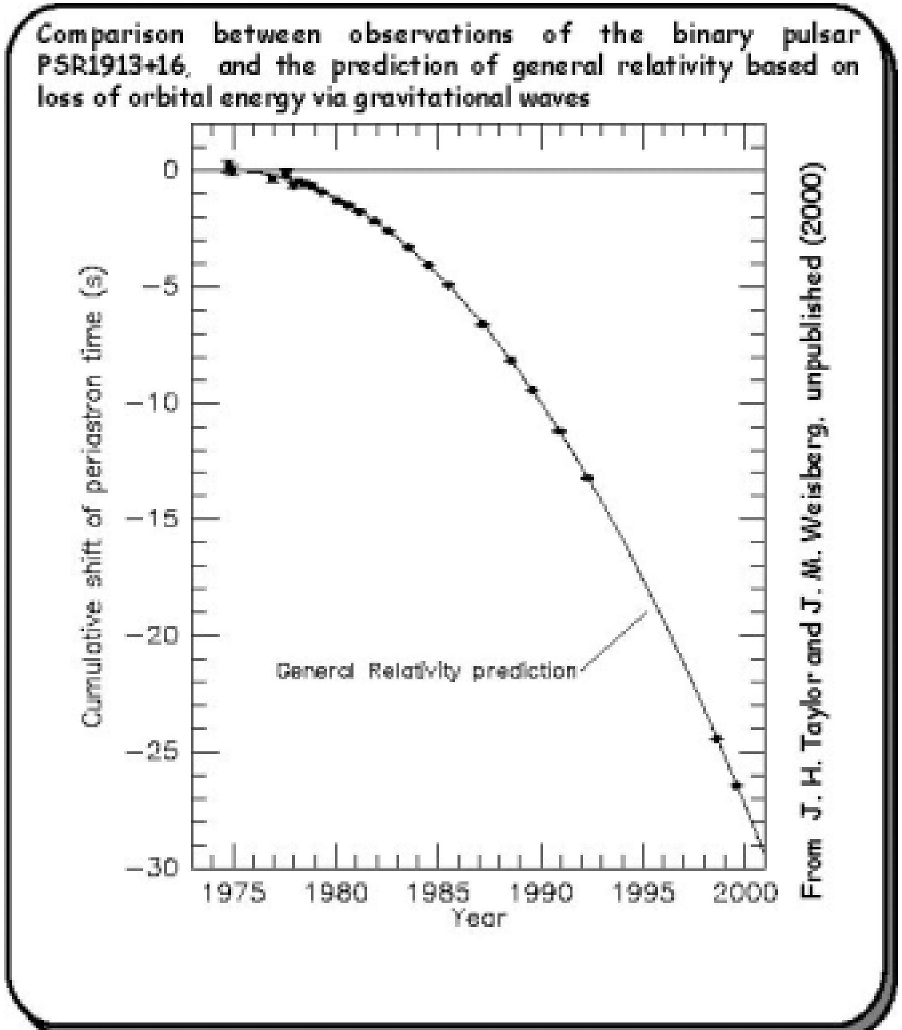
PSR 1913 + 16 -- Timing of pulsars



Neutron Binary System
- **separated by 10^6 miles**
- **$m_1 = 1.4m_{\odot}$; $m_2 = 1.36m_{\odot}$; $\varepsilon = 0.617$**

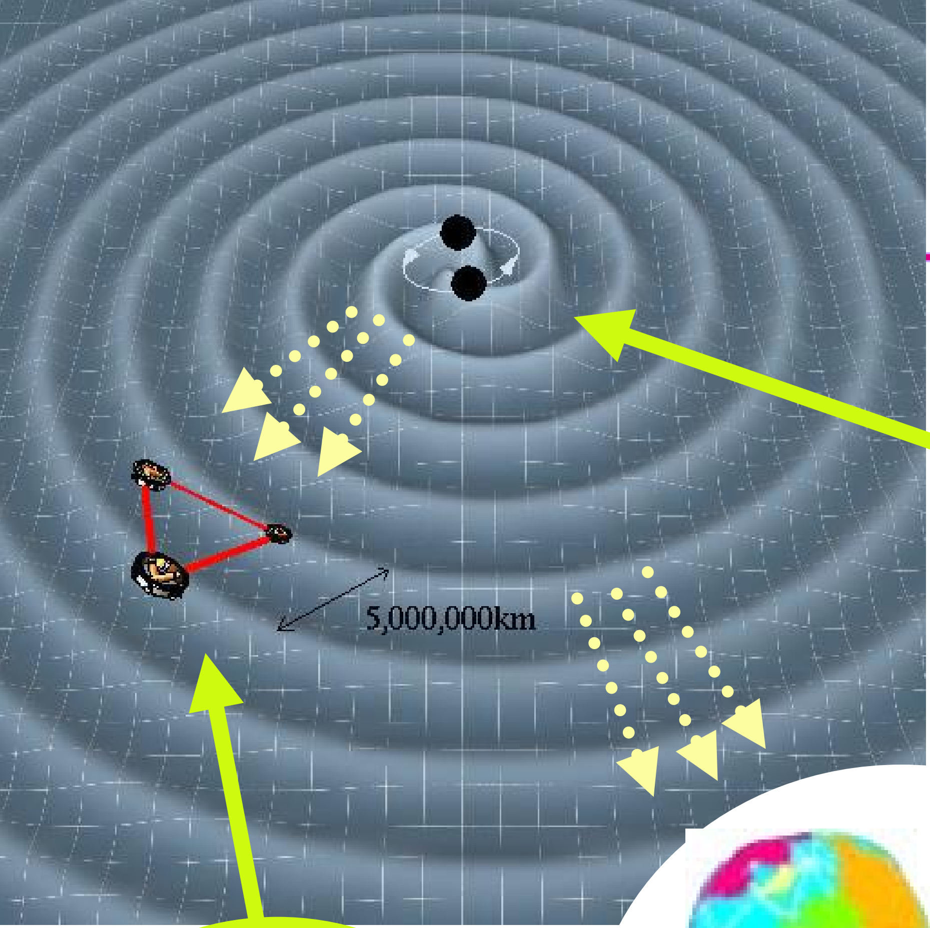
Prediction from general relativity
- **spiral in by 3 mm/orbit**
- **rate of change orbital period**

Emission of gravitational waves



Direct Detection

astrophysical sources



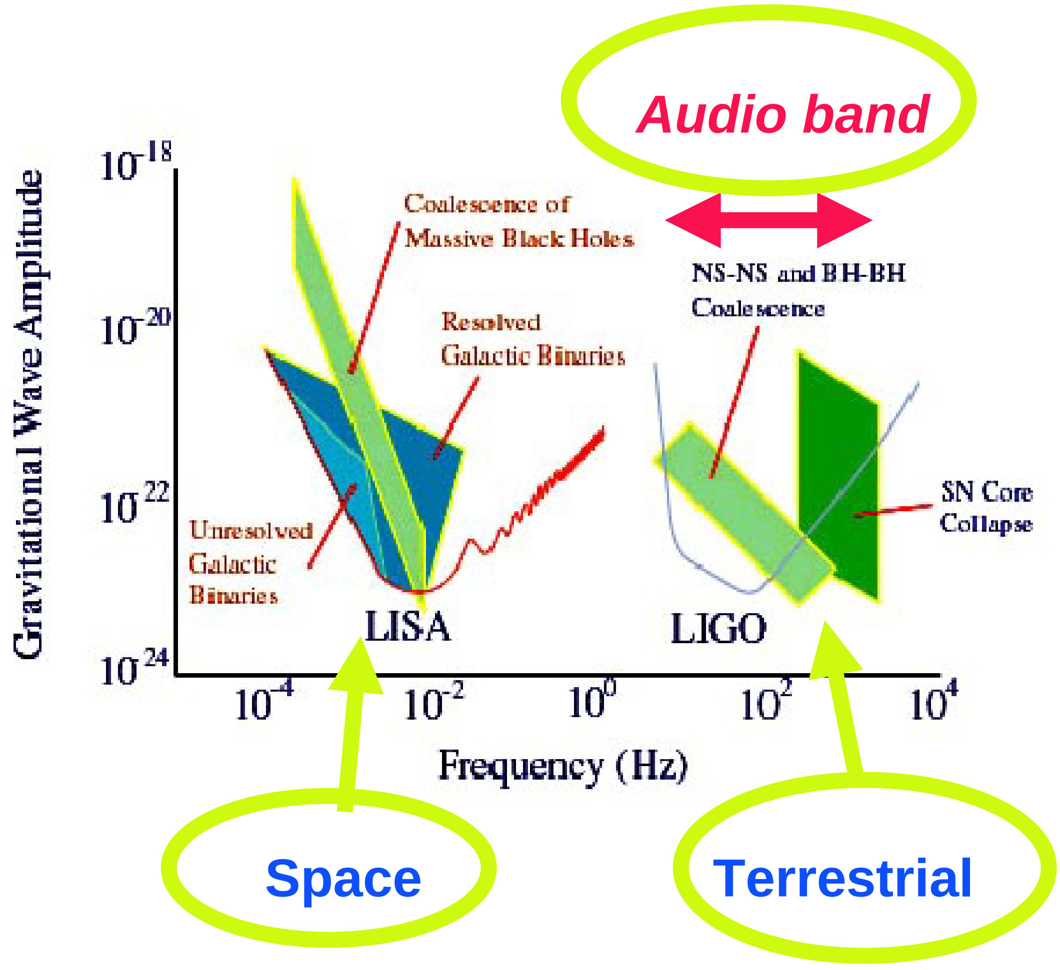
Gravitational Wave Astrophysical Source

Detectors in space
LISA

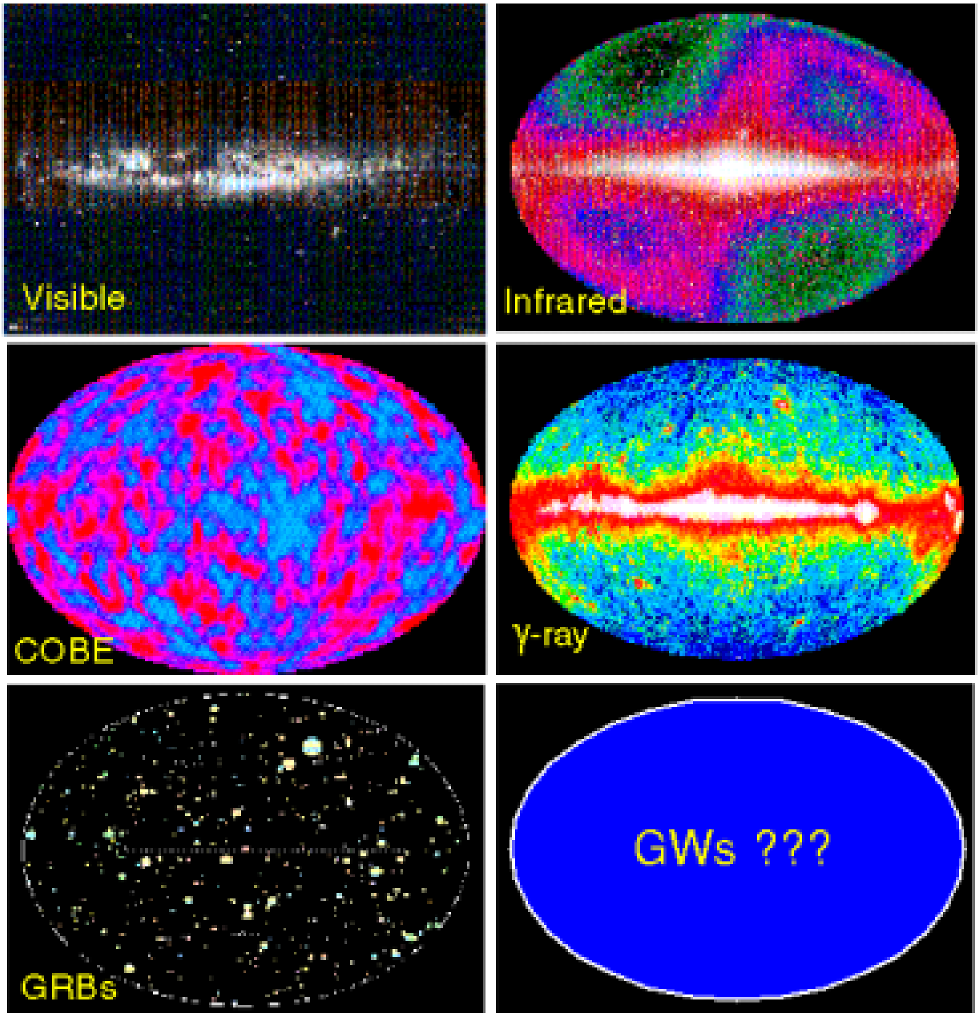
Terrestrial detectors
LIGO, TAMA, Virgo,AIGO

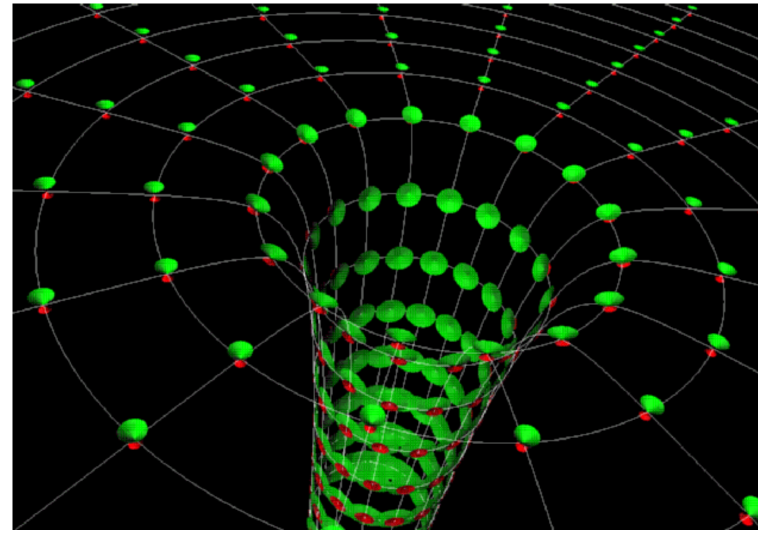
Astrophysics Sources
frequency range

- **EM waves are studied over ~20 orders of magnitude**
  - » **(ULF radio –> HE γ-rays)**
- **Gravitational Waves over ~10 orders of magnitude**
  - » **(terrestrial + space)**



A New Window on the Universe



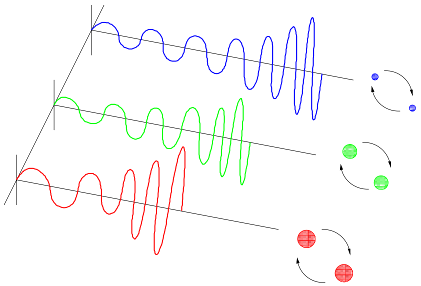


Gravitational Waves will provide a new way to view the dynamics of the Universe

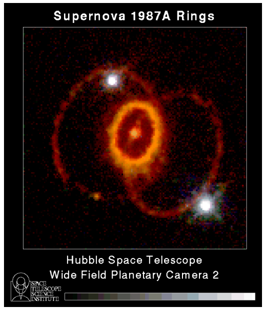
LIGO

Astrophysical Sources *signatures*

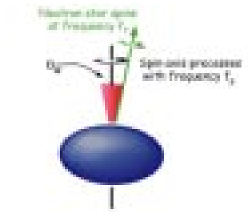
- **Compact binary inspiral:** ***"chirps"***
  - » **NS-NS waveforms are well described**
  - » **BH-BH need better waveforms**
  - » **search technique: matched templates**



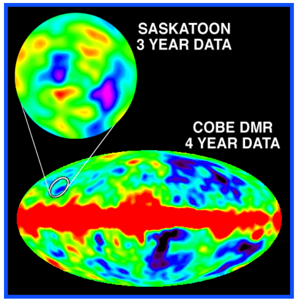
- **Supernovae / GRBs:** ***"bursts"***
  - » **burst signals in coincidence with signals in electromagnetic radiation**
  - » **prompt alarm (~ one hour) with neutrino detectors**



- **Pulsars in our galaxy:** ***"periodic"***
  - » **search for observed neutron stars (frequency, doppler shift)**
  - » **all sky search (computing challenge)**
  - » **r-modes**



- **Cosmological Signals** ***"stochastic background"***



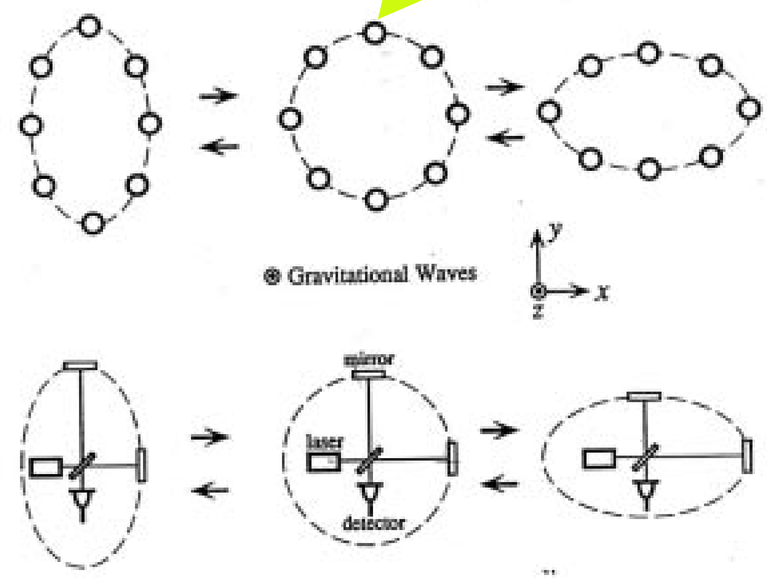
Joseph Weber 1919-2000

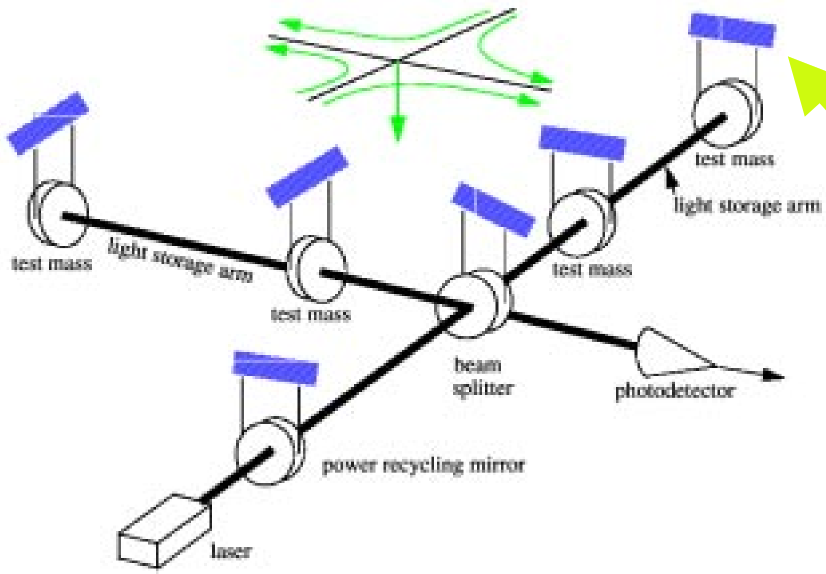
Interferometers

terrestrial

free masses

International network (LIGO, Virgo, GEO, TAMA, AIGO) of suspended mass Michelson-type interferometers on earth's surface detect distant astrophysical sources

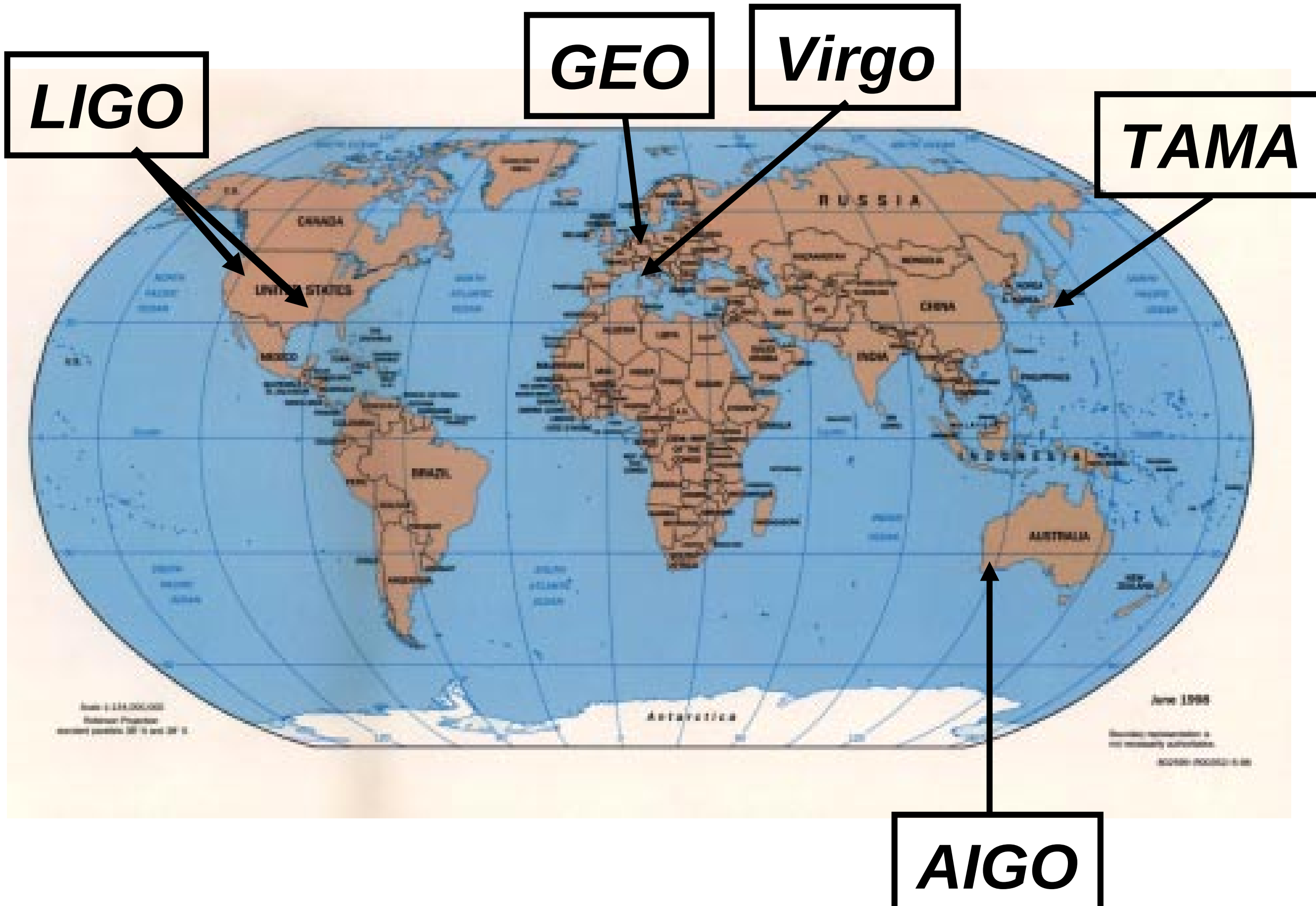




suspended test masses

Interferometers
international network

Simultaneously detect signal (within msec)



detection confidence

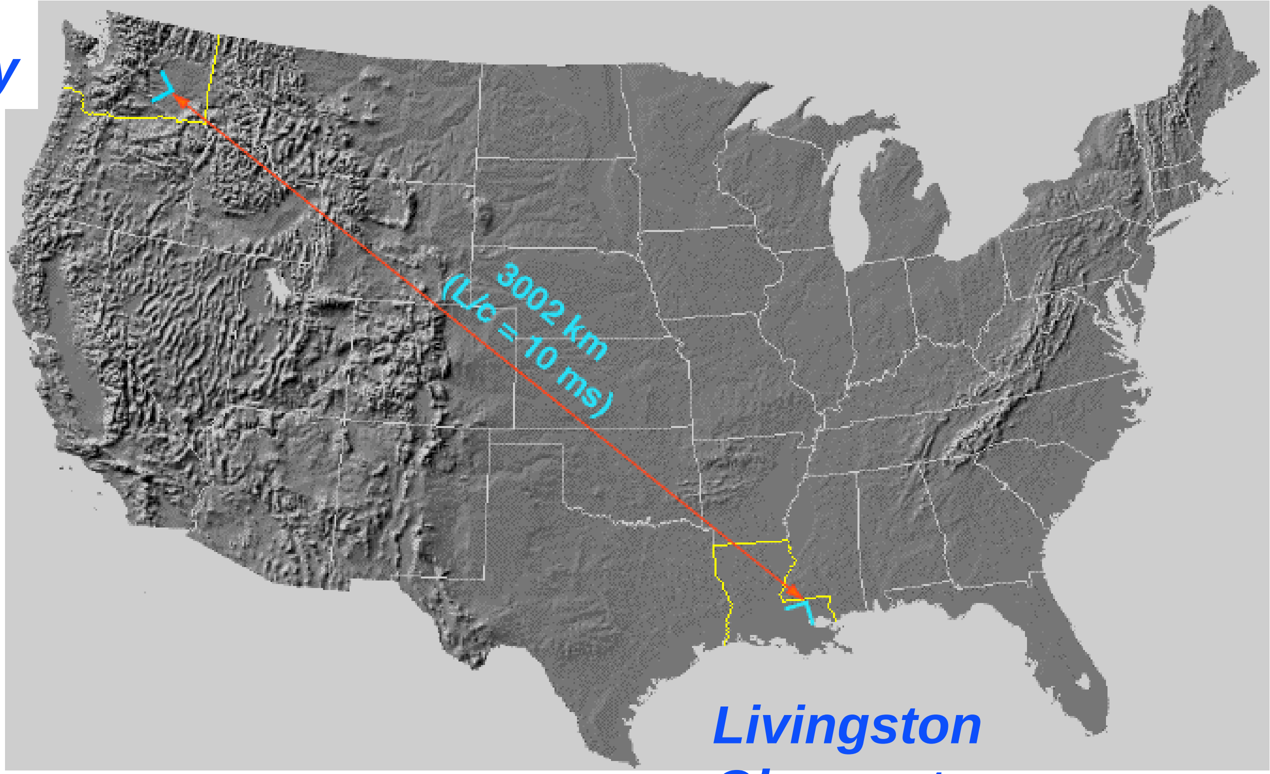
locate the sources

decompose the polarization of gravitational waves

The Laboratory Sites

<u>*Laser Interferometer Gravitational-wave Observatory*</u> *(LIGO)*

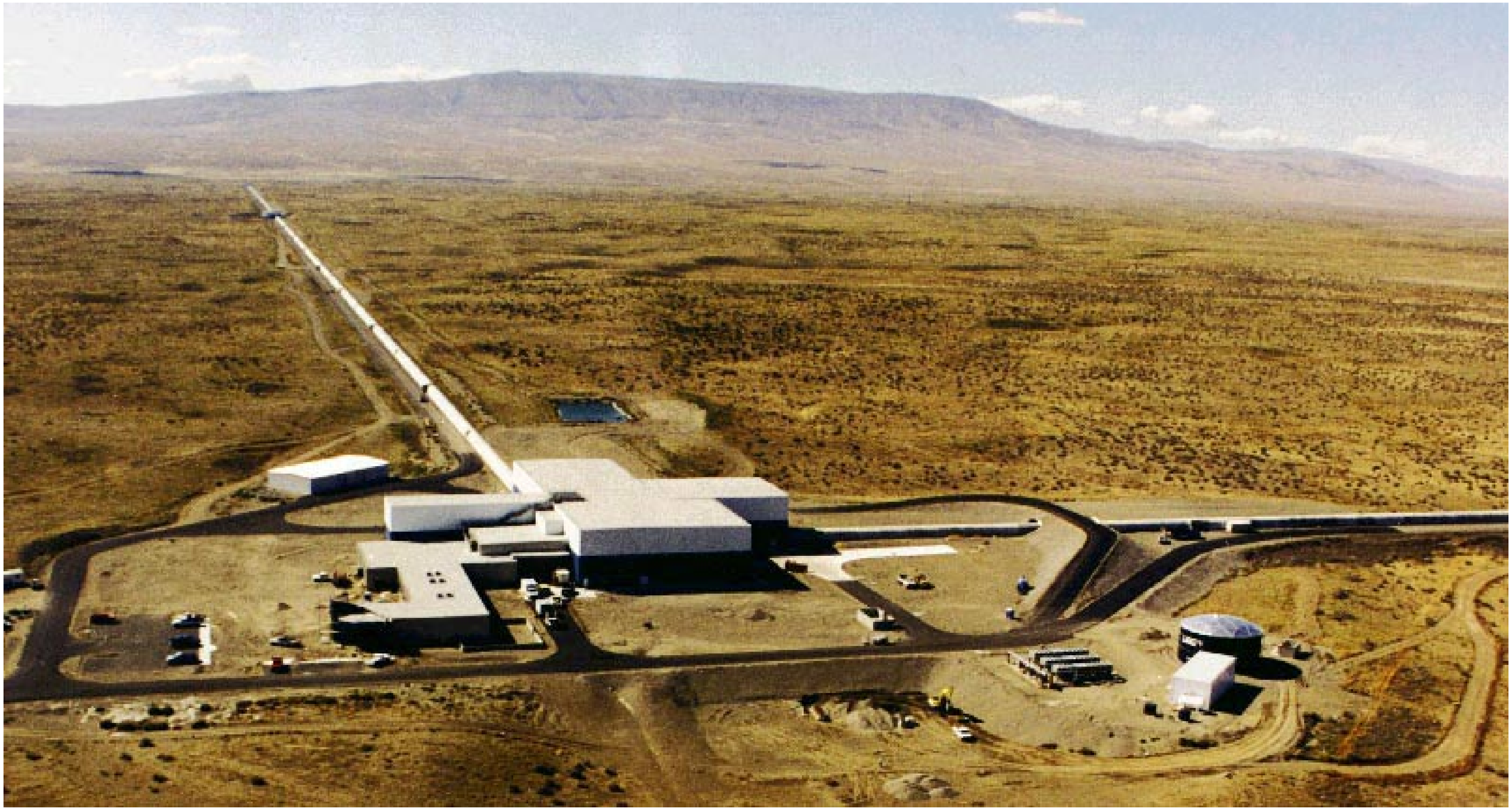
Hanford Observatory



Livingston Observatory

LIGO
Livingston Observatory

LIGO
Hanford Observatory



LIGO
beam tube

- LIGO beam tube under construction in January 1998
- 65 ft spiral welded sections
- girth welded in portable clean room in the field

1.2 m diameter - 3mm stainless
50 km of weld

NO LEAKS !!

LIGO
vacuum equipment

LIGO Optic

Substrates: SiO_2

25 cm Diameter, 10 cm thick

Homogeneity < 5 x 10^{-7}

Internal mode Q's > 2 x 10^6

Polishing

Surface uniformity < 1 nm rms

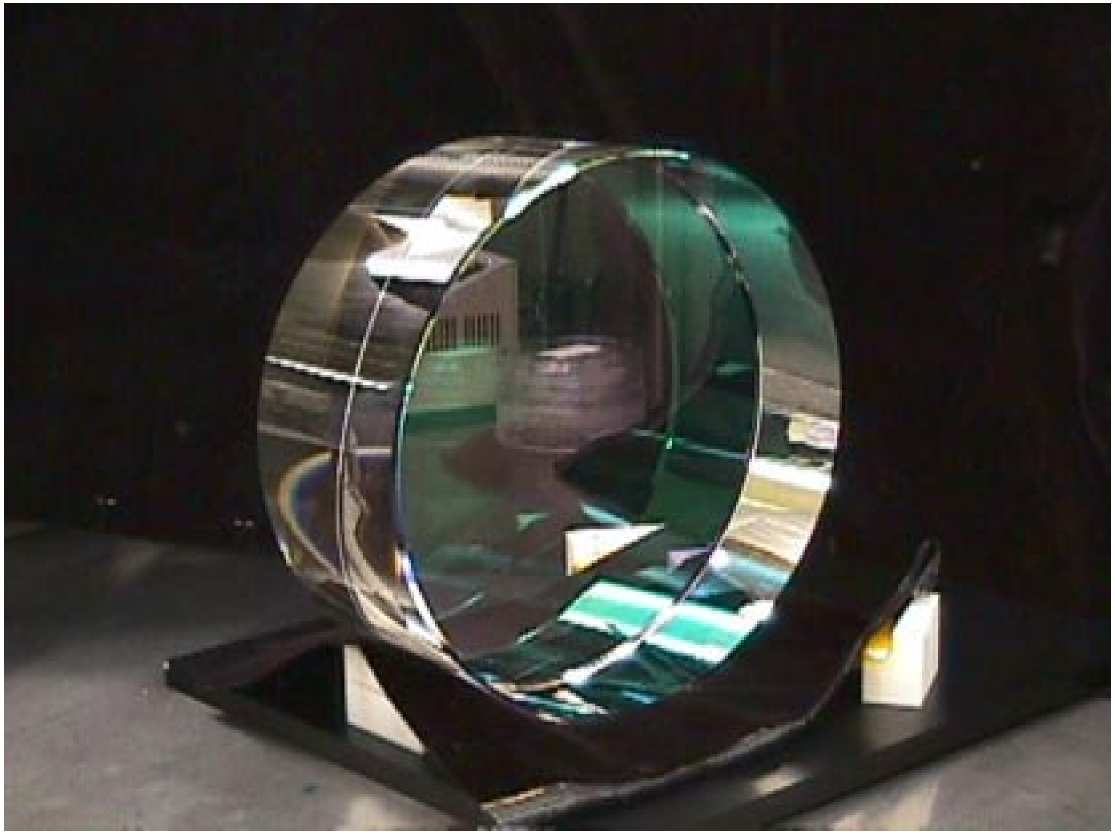
Radii of curvature matched < 3%

Coating

Scatter < 50 ppm

Absorption < 2 ppm

Uniformity <10^{-3}



LIGO

Core Optics

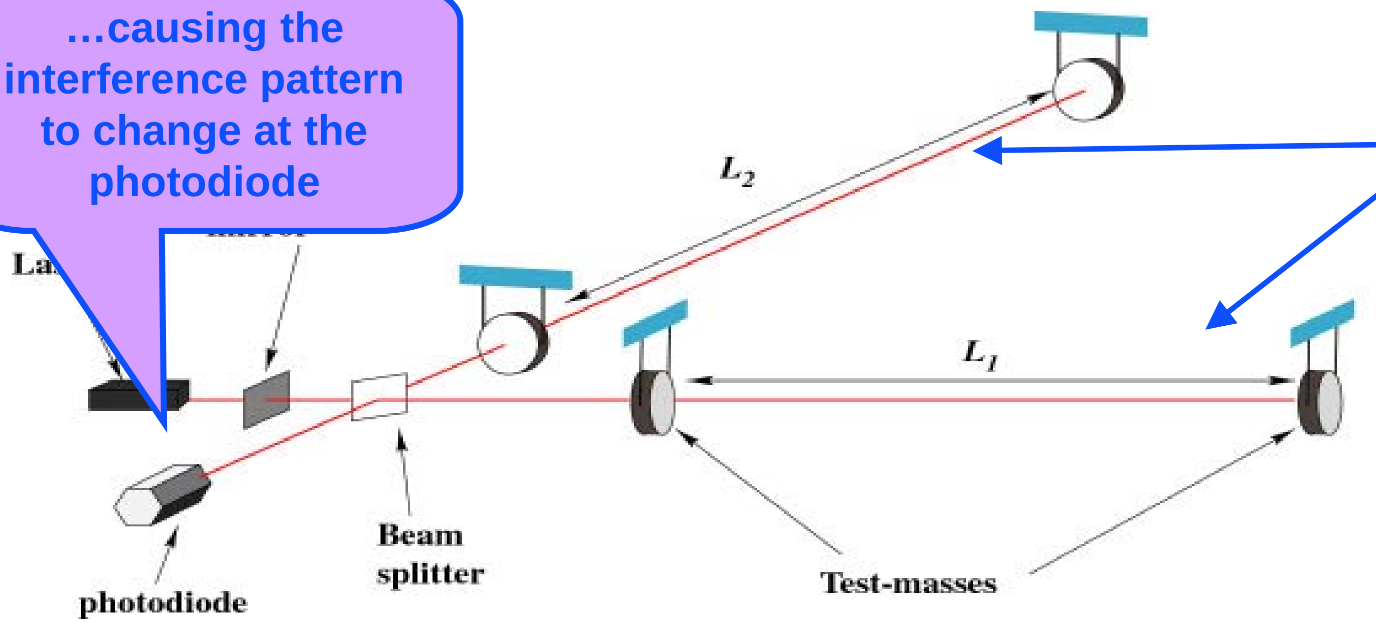
installation and alignment

Suspended Mass Interferometer
the concept

- **An interferometric gravitational wave detector**
  - » **A laser is used to measure the relative lengths of two orthogonal cavities (or arms)**

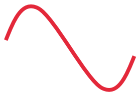
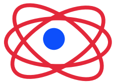
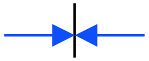
- **Arms in LIGO are 4km**
  - » **Current technology then allows one to measure *$h = \delta L/L \sim 10^{-21}$* which turns out to be an interesting target**

…causing the interference pattern to change at the photodiode



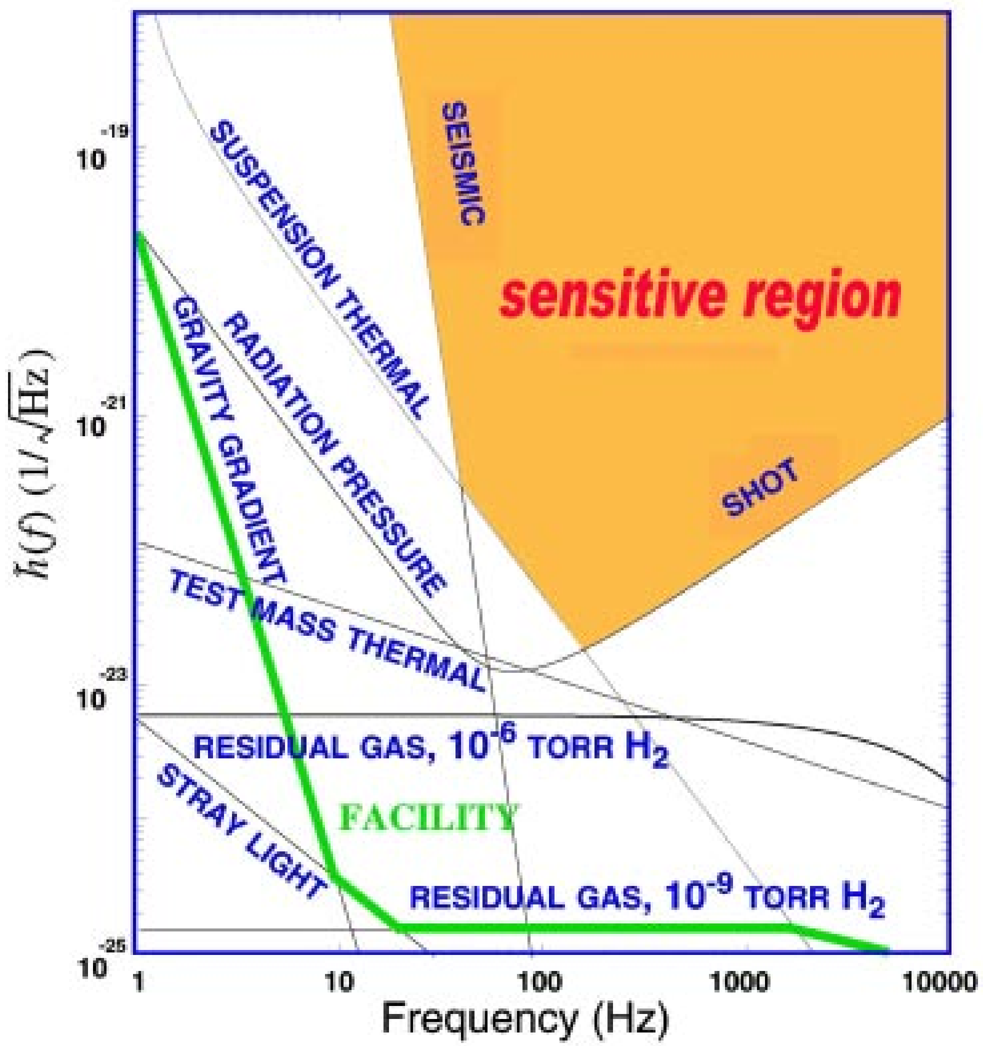
As a wave passes, the arm lengths change in different ways….

How Small is 10^{-18} Meter?

		One meter, about 40 inches
÷10,000		*Human hair, about 100 microns*
÷100		*Wavelength of light, about 1 micron*
÷10,000		*Atomic diameter, 10^{-10} meter*
÷100,000		*Nuclear diameter, 10^{-15} meter*
÷1,000		*LIGO sensitivity, 10^{-18} meter*

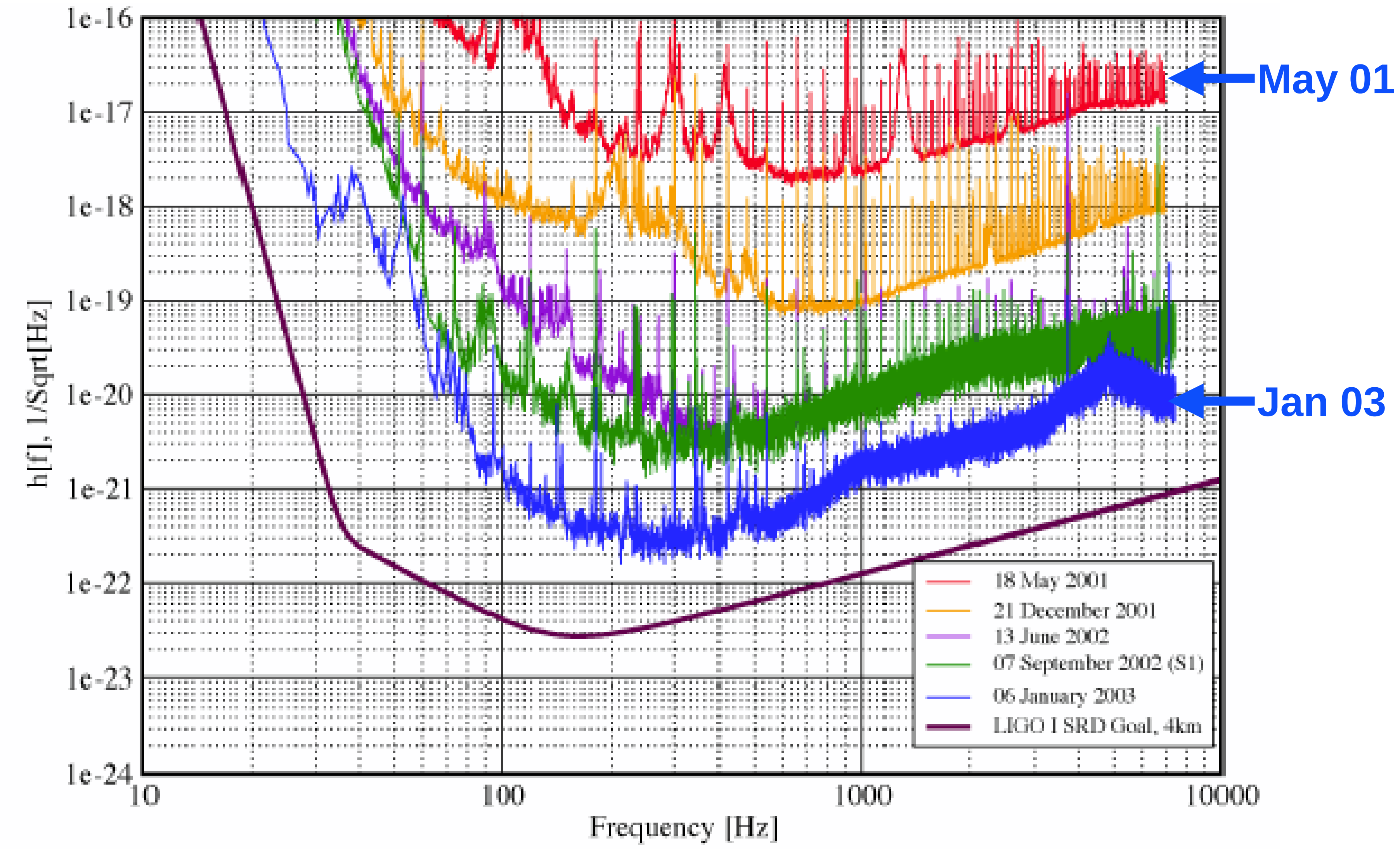
What Limits Sensitivity of Interferometers?

- Seismic noise & vibration limit at low frequencies
- Atomic vibrations (Thermal Noise) inside components limit at mid frequencies
- Quantum nature of light (Shot Noise) limits at high frequencies
- Myriad details of the lasers, electronics, etc., can make problems above these levels



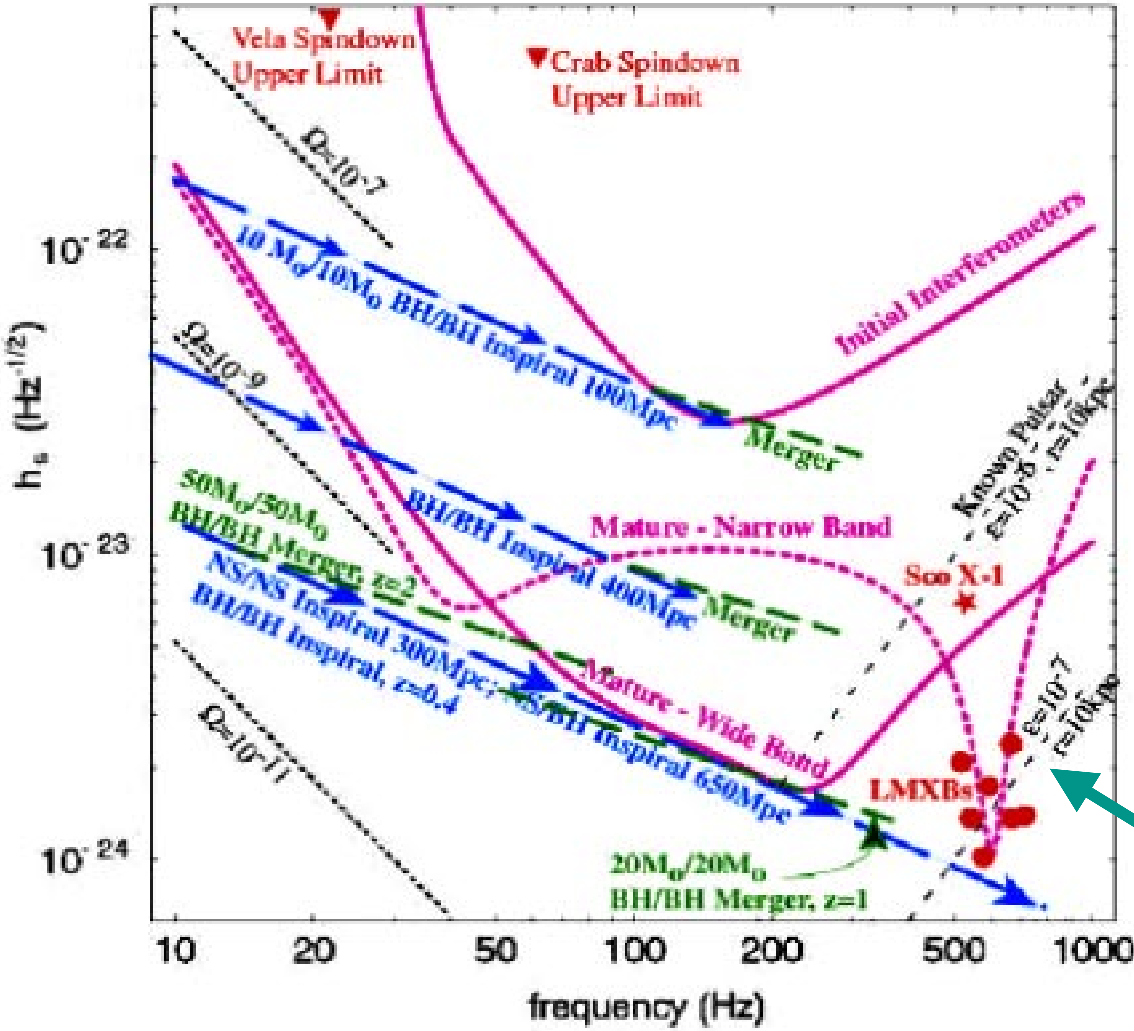
LIGO Sensitivity

Livingston 4km Interferometer



Advanced LIGO

2007 +



Enhanced Systems

- laser
- suspension
- seismic isolation
- test mass

Improvement factor in rate $\sim 10^4$

+

narrow band optical configuration

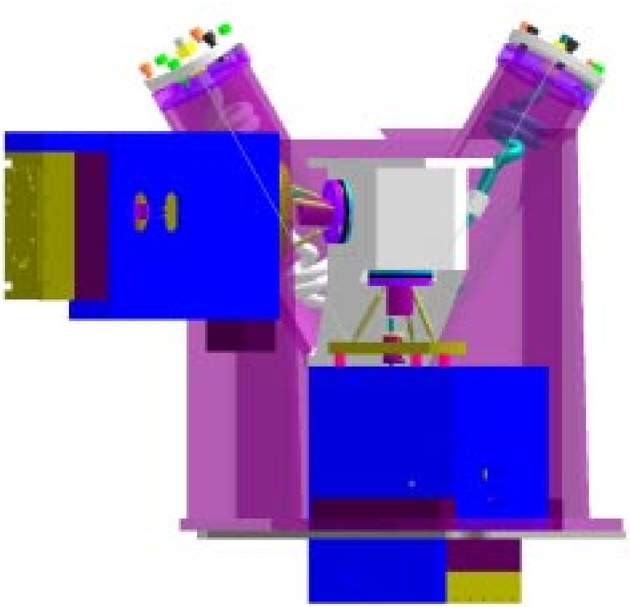
Advanced LIGO

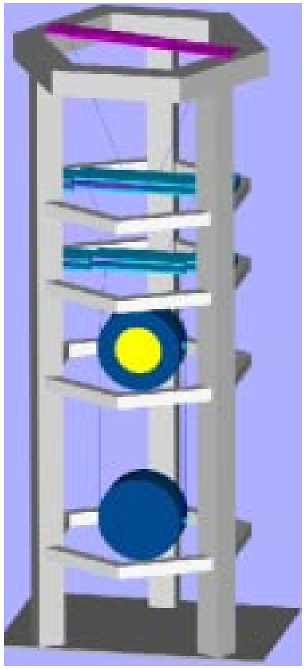
improved subsystems

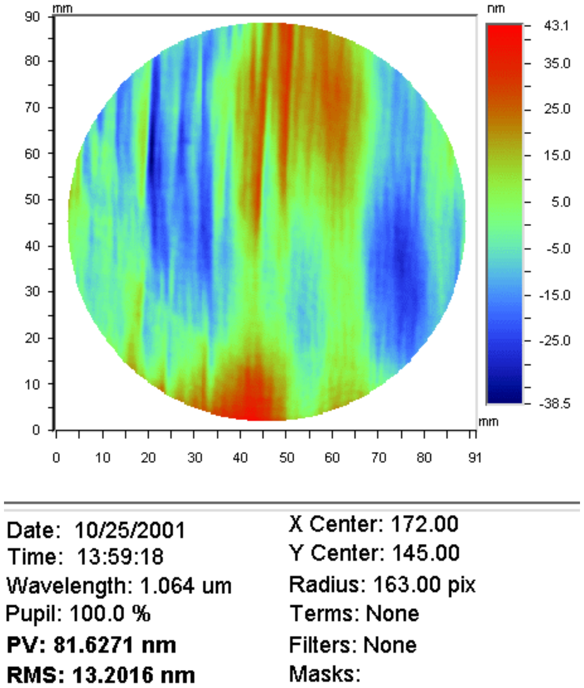
Multiple Suspensions

Active Seismic

Sapphire Optics

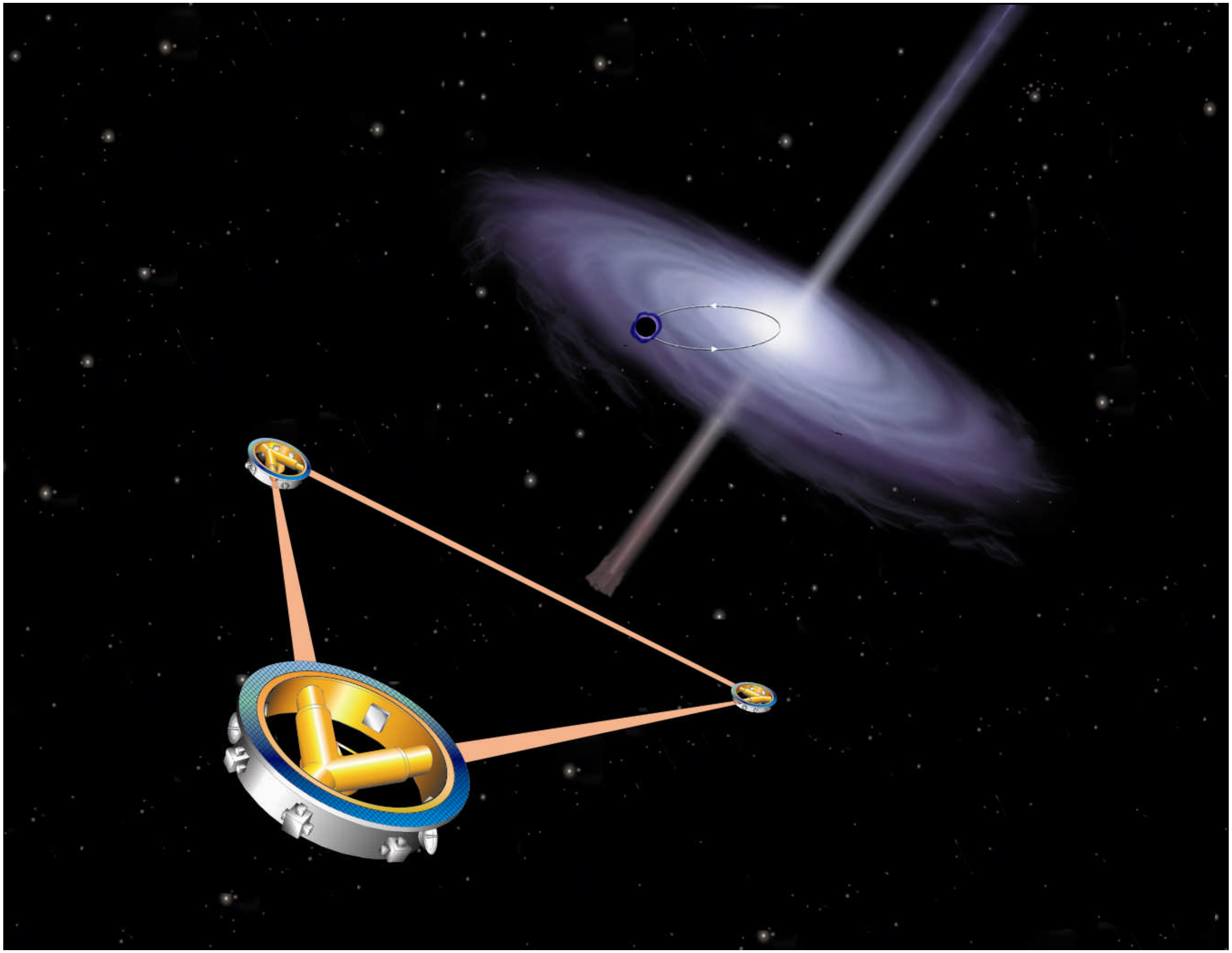






Higher Power Laser

Mission Concept



Optical System